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Far-ultraviolet electrographic images, covering the wave and 1230 to 1600 %, were obtained of the Large Magellan mission with the NRL Far-Ultraviolet Camera/Spectrograp images were analyzed as described in the "S201 Far-Ultrav Magellanic Cloud" (NRL Report 8206), and the "Revised	ic Cloud during the Apollo 16 oh (Experiment S201). The violet Atlas of the Large
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20. ABSTRACT (Continued)

Objects" (NRL Report 8487). The listing of far-UV objects in the LMC provided here is improved in quantitative accuracy and completeness compared to that originally provided in NRL Report 8206. These improvements include more accurate estimates of interstellar extinction corrections, comparisons with more sensitive ground-based Hage emission measurements, and calibration of the S201 photometry by comparison with OAO-2 photometry of objects in common. The listing presents ultraviolet brightnesses, measured on one or more of four frames (two 1050-1600 A and two 1250-1600 A exposures), for 473 objects or groupings of objects in the LMC. Also listed, where available, are Ha brightnesses of associated nebulosities, and values of hydrogen index (defined as ratio of Ha brightness to far-UV brightness, both corrected for interstellar extinction).

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REVISED LISTING - S201 FAR-ULTRAVIOLET ATLAS OF THE LARGE MAGELLANIC CLOUD

I. INTRODUCTION

Far-ultraviolet imagery of the Large Magellanic Cloud was obtained with the Naval Research Laboratory's Far Ultraviolet Camera (Experiment S201) during the Apollo-16 mission, 21-23 April 1972. This imagery covered two wavelength ranges, 1050-1600 Å and 1250-1600 Å, with a limiting angular resolution of about 3 arc min.

The analysis of the Large Magellanic Cloud imagery was discussed in NRL Report 8206, S201 Far Ultraviolet Atlas of the Large Magellanic Cloud (July 12, 1978) and in a paper published in The Astrophysical Journal ("Distributions of Hot Stars and Hydrogen in the Large Magellanic Cloud," 15 September 1981). The purpose of the present Memorandum Report is to present a revised listing of the individual objects or regions in the LMC which were detected and measured, similar to that in the original Atlas but with the following improvements and additions:

- 1. The extinction correction for objects observed in the LMC has been modified based on recent studies of the LMC interstellar extinction using the International Ultraviolet Explorer (IUE) satellite. Also, extinction at the hydrogen Balmer- α (6563 Å) wavelength has been taken into account in the derivation of Hydrogen Index values.
- 2. More recent $H\alpha$ observations by Davies, Elliott, and Meaburn (1) have been included in the Hydrogen Index derivations.
- 3. The S201 photometry has been compared with OAO-2 photometry of stars in common to place the S201 measurements on an absolute scale of ultraviolet brightness.

II. DATA AND ANALYSIS

The far-ultraviolet images of the Large Magellanic Cloud are qualitatively useful for determining the spatial distributions of early-type stars in the LMC without confusion by images of the far more numerous cooler stars (almost all stars detected in the S201 imagery are of spectral type earlier than A2; i.e., with effective temperatures above 9000 K). The distribution of hot stars differs considerably from the general stellar population

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distribution as revealed by visual imagery; the short (3 min) 1250-1600 Å exposure shows the previously known OB associations and clusters, whereas the longest (30 min) 1250-1600 Å exposure shows the general distribution of hot stars, most of which are less luminous than those in the associations. Comparison of the UV imagery with Hα and blue imagery (1) indicates that, for the most part, the extended nebulosities in the LMC (many of which are considerably larger than the S201 resolution limit) are not conspicuous in the far-ultraviolet. This is also indicated by IUE observations of the 30 Doradus nebula (2) and of local galactic H II regions. Thus, the observed far-UV is presumed to be either direct starlight or starlight scattered by dust in close proximity to the stars.

Quantitative analysis of the imagery is, to some extent, complicated by the effects of interstellar extinction, correction for which is particularly uncertain in the LMC because of incomplete knowledge of E(B-V) and of the extinction vs. wavelength in the LMC. It is known from ANS and IUE observations that the LMC extinction law is considerably different from that applicable in the local regions of our galaxy and shows large variation with position in the LMC (3).

Ultraviolet Photometry

The procedures used for the reduction and processing of the S201 electrographic imagery have been presented in detail in our Far Ultraviolet Atlas of the Large Magellanic Cloud (NRL Report 8206) and in the Revised S201 Catalog of Far Ultraviolet Objects (NRL Report 8487). In summary, for any identifiable image, the integrated intensity is proportional to the density volume $V = \sum (d_L - b_L)$. Here, d_L and b_L are the optical densities D_L and B_L (as measured by the PDS microdensitometer used to scan the films) times 100, of each pixel in the image and in background areas near (but outside) the image, respectively; the sum is over all pixels detectably above the adopted background. The subscript L indicates that the densities have been corrected for nonlinearities of the emulsion and microdensitometer. The density volume can then be related to ultraviolet brightness by reference to preflight calibrations of the instrument and/or comparison of observations of objects in common with other photometrically calibrated observations, such as those of OAO-2 (4). We have determined,

through comparison of our preflight calibration predictions with OAO-2 measurements by Code et al. (4) that the absolute sensitivity of the S2O1 camera was probably a factor of 1.5 (0.45 stellar magnitudes) less, at the time of the observations, than predicted by our preflight calibrations.

In the LMC, determination of the UV brightnesses of individual objects is difficult, because of the limited resolution of our imagery and because of the multitude of field stars against which an individual object must be observed. This makes determination of the true background which should be subtracted from the measured density, in determinations of the density volumes, very uncertain. However, contour plots such as those in the Atlas (NRL Report 8206) give useful measurements of the ultraviolet brightness distribution over the face of the IMC, which are significant to studies of the interstellar medium in the LMC (photoionization and photodissociation equilibria of many interstellar species are largely controlled by the stellar ultraviolet radiation field longward of 912 Å) and which, in conjunction with other determinations of stellar spectral type or effective temperature, provide indications of the distribution of dust extinction over the LMC. In areas where the individual stellar contributions cannot be resolved, the local surface brightness (above sky background outside the LMC) may be useful for some purposes. Based on our preflight calibrations, a density above background of 0.1 D corresponds to an intensity of 1.89×10^6 photons/cm² sec sterad at the effective wavelength (1400 Å) of the camera with CaF2 corrector (wavelength range 1250-1600 Å). For a flat continuum extending over the camera effective passband of 250 Å, this corresponds to 7.56 x 10^3 photons/cm² sec $^{\text{A}}$ sterad (1.07 x $^{\text{10}}$ erg/cm² sec $^{\text{A}}$ sterad).

We obtained a measure of the total UV brightness of the LMC in the 1050-1600 Å and 1250-1600 Å bands by summing the densities above sky background of all pixels in the LMC region of each frame. The contributions of seven SAO stars were subtracted. The total brightness of the LMC (based on our preflight calibrations) in the 1250-1600 Å wavelength range ($\lambda_{\rm eff}$ = 1400 Å) is 220 photons/cm² sec Å or F₁₄₀₀ = 3.12 x 10⁻⁹ ergs/cm² sec Å. This corresponds to a UV magnitude, in the system of Code et al. (4), of m₁₄₀₀ = 0.23. In the 1050-1600 Å range ($\lambda_{\rm eff}$ = 1300 Å) the corresponding UV magnitude is m₁₃₀₀ = 0.13.

Averaged over the apparent angular size of the LMC on our image (about 6° diameter, or 9 x 10^{-3} sterad) the mean surface brightness is S_{1400} = 2.4×10^4 photons/cm² sec Å sterad (3.4 x 10^{-7} ergs/cm² sec Å sterad), and S_{1300} = 2.5×10^4 photons/cm² sec Å sterad (3.8 x 10^{-7} ergs/cm² Å sterad).

These measurements include both direct and dust-scattered starlight (we assume that nebular emission lines make a negligible contribution to the total UV brightness). As mentioned earlier, use of the OAO-2 photometry as a reference standard will increase the above intensity by a factor of 1.5. Except for a minor correction due to interstellar extinction within our galaxy in the line of sight to the IMC, this gives an indication of the local stellar radiation field, on the average, within the LMC. The average surface brightness at 1400 Å corresponds to a radiation density of $U_{1400} = \frac{4\pi}{c} \, S_{1400} = 1.4 \times 10^{-16} \, \text{ergs/cm}^3$ Å. This may be compared with estimates of the radiation field within our own galaxy of about $10^{-16} \, \text{ergs/cm}^3$ Å at 1400 Å by Witt and Johnson (5) and about half this value predicted by Henry (6).

In the Revised Listing (Appendix B), we present our best estimates of the LTV brightnesses of individual objects (actually, individual brightness peaks in our imagery) above the local background (defined individually for each brightness peak). The net density volume divided by exposure, V/E (uncorrected for interstellar extinction) is a direct measure of the ultraviolet brightness, as discussed in the next section. These V/E values can be converted into absolute ultraviolet magnitudes, as discussed in the Revised S201 Catalog, using the relationships

$$m_L = 14.13 - 2.512 \log (V/E)_L$$

$$m_C = 13.18 - 2.512 lg (V/E)_C$$

Here, subscript L indicates exposures in the 1050-1600 Å wavelength range (LiF corrector, effective wavelength 1300 Å), which in the LMC included Frames 124 (1 min exposure) and 125 (3 min exposure). Subscript C designates exposures in the 1250-1600 Å range (CaF₂ corrector, effective wavelength 1400 Å), which includes Frames 129 (10 min exposure) and 130 (30 min exposure).

Hydrogen Index

In the Atlas we derived a "hydrogen index" (hereafter H Ind) as the ratio of H α flux, HA, to far-UV flux, UF (corrected for dust extinction),

at over 100 places in the LMC. This index was first presented as a rough measure of the hydrogen near hot stars or star groups detected on our far-UV images. That is, if the ionizing extreme-UV (λ < 912 Å) flux is assumed roughly proportional to the far-UV flux, then the intensity of H α emission is related to the local hydrogen density. Here, we present a revised determination of H Ind and its variation over the LMC, using a more recent determination of the LMC extinction law, allowing for extinction at H α as well as in the UV, and utilizing additional data on the H α brightness distribution in the LMC.

The far-UV flux values are proportional to the measured density volume, V (corrected for nonlinearities and background) divided by the exposure time, F, in minutes. As shown in the <u>Revised S201 Catalog of Far-UV Objects (NRL Report 8487)</u>, a density-volume

$$V = 0.037 \text{ n}$$
 (1)

where n is the number of photoelectrons forming the far-UV image. Thus,

$$V/E = 6.17 \times 10^{-4} \text{ n per sec}$$
 (2)

where E is the exposure time in min, and n/sec is related to the photons arriving each sec from the object. The detection efficiency (photo-electrons per photon, based on preflight calibrations) of the S201 Camera in the imaging mode averages 0.05 over the range 1050-1600 Å with the LiF corrector, and 0.04 over the range 1250-1600 Å with the CaF₂ corrector. Hence, the photon flux in these wavelengths is

$$N_L = n_L/0.05(30.0) = 1.08 \times 10^3 \text{ (V}_L/\text{E) photons/sec cm}^2 \text{ for } 1300 \text{ Å} + 250 \text{ Å},$$
 (3)

 $N_C = n_C/0.04(30.0) = 1.35 \times 10^3 \ (V_C/E)$ photons/sec cm² for 1400 Å \pm 150 Å, (4) where 30.0 cm² is the aperture area of the S201 camera. Since these photons each carry 1.52 x 10^{-11} erg and 1.42 x 10^{-11} erg respectively, the far-UV flux is

$$F_L = 1.64 \times 10^{-8} \ (V_L/E) \ erg \ sec^{-1} \ cm^{-2}$$
 (5)

$$F_C = 1.92 \times 10^{-8} \text{ (V}_C/E) \text{ erg sec}^{-1} \text{ cm}^{-2}$$
. (6)

These were corrected for interstellar extinction, based on previous estimates (7) of the visual reddening (RE = E(B-V)). In order to estimate reddening for all our measurements of V/E, for which specific values of RE were not available, we plotted Lucke's (7) RE values and sketched in

contour lines (see Fig. 1). Although Lucke's 81 measured values are good to + 0.05, corresponding to + 16 to + 17% in corrected ultraviolet flux, UF, there is inevitably some uncertainty in the interpolated values of RE, due to small scale variations in the extinction at a given distance, and the uncertainty in distance to an object along the line of sight. The stellar associations for which Lucke determined RE may lie in front of or behind far-UV objects with nearly the same celestial coordinates. However, it is highly likely that an LH cluster and an associated Fenize nebula are in close 3-dimensional proximity.

In the <u>Atlas</u>, we used the "average" galactic interstellar extinction curve of Bless and Savage (8). However, measurements with the ANS satellite (9,10) in the °C Doradus region, and with IUE (3) there and elsewhere in the LMC indicate a higher ratio of far-UV extinction to E(B-V) in the LMC than is typical in the local region of our galaxy (see Figure 2). Using the extinction curve of Ref. (3) with $A_{\lambda} = 3$ E(B-V) + E(λ -V), we have, for effective wavelengths of 1300 Å (LiF corrector) and 1400 Å (CaF₂ corrector), E(1300-V)/E(B-V) = 8.97 and E(1400-V)/E(B-V) = 7.09. Therefore, the ultraviolet fluxes corrected for reddening are

$$UF_{L} = F_{L} 10^{4.8 \text{ RE}}$$
 (7)

$$UF_{C} = F_{C} 10^{4.0} RE$$
 (8)

As expected, UF_L values for an object are generally larger than the UF_C values because of the wider bandpass and larger extinction correction at the effective wavelength of 1300 Å. The scatter in the LMC extinction curve of Nandy et al. (3) is about 0.2 mag. The extinction correction at H α is assumed to be A_{6563} = 2.5 RE; hence the corrected H α flux is UHA = HA·10^{RE}, approximately, where HA is the H α flux as measured by Henize et al. (11,12) in units of 10^{-4} erg/cm² sec sterad. The HA values given here are often summed for several close H II regions that could not be separately resolved on our S201 photos. For instance, N180A-C means the summed flux from N180A, N180B, and N180C. In order to get a single hydrogen index representing all measurements of a given object, we averaged the values for two ILi frames with 1/2 the values for two ICa frames:

$$H Ind_{L} = UHA/UF_{L}$$
 (9)

$$H Ind_{C} = UHA/UF_{C}$$
 (10)

H Ind = (H Ind_{L1} + H Ind_{L2} +
$$1/2$$
 H Ind_{C1} + $1/2$ H Ind_{C2})/4 (11)

The major errors in V/E, UF, and H Ind are due to uncertainty in background, b. As can be seen from the isodensity contour plots in the Atlas, many of the objects measured are in regions where the background density is changing. The local background was estimated on mosaics of d, taking the first minimum in d in each of four directions from the peak density, along +x, +y, -x, and -y, and averaging these to get b. The background is high and posed the most difficulties on the 3-min ILi exposure, frame A125.

The HA values are probably good to \pm 10%, although values near zero are subject to larger percentage errors. In fact, DEM, in a careful survey of a 5-hour exposure with the SRC 48-inch Schmidt camera using an interference filter with 100 Å bandpass centered on H α and [NII], found the faint Henize H II regions much larger, and detected 100 more, most of them fainter than Henize's limit. They give no quantitative measurements of brightness, but use the steps vf (very faint), f (faint), fb (fairly bright), b (bright), and vb (very bright). We calibrated this scale against HA by assigning the numbers vf = 1, f = 2, fb = 3, b = 5, vb = 10, and multiplying by the dimensions given in arc-min. For instance, a faint (f) nebula of size 3.5' x 2' has a brightness (arc-min)² of 2 x 3.5 x 2 = 14. Fig. 3 is a plot of these values against HA for 64 cases where the DEM dimensions are roughly the same as Henize's. To a fairly good approximation,

PEM brightness $(arc-min)^2 = 3 \text{ HA}$.

(12)

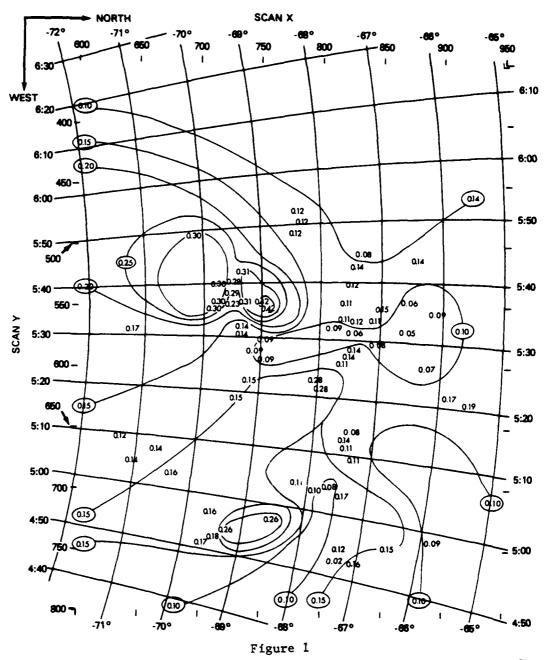
Using this calibration, we could fill in 227 H II regions at positions in the LMC where we had measured far-UV flux, leaving out only 19 DEM objects of the total of 356. (These positions were all searched on our mosaics.)

In the Revised Listing (Appendix B), we list density volumes for 473 objects or regions in the LMC. We also list values of UF, defined here as simply the density volumes corrected for extinction as per Equations 7 and 8. (True UF values, in ergs/sec cm², can be obtained by multiplying by the factors 1.64×10^{-8} for F_L and 1.92×10^{-8} for F_C , respectively.) Likewise, the H Ind values are the corrected density volumes divided by UHA. Figure 4 is a contour plot of H Ind (times 100), individual values of which are given in the Listing.

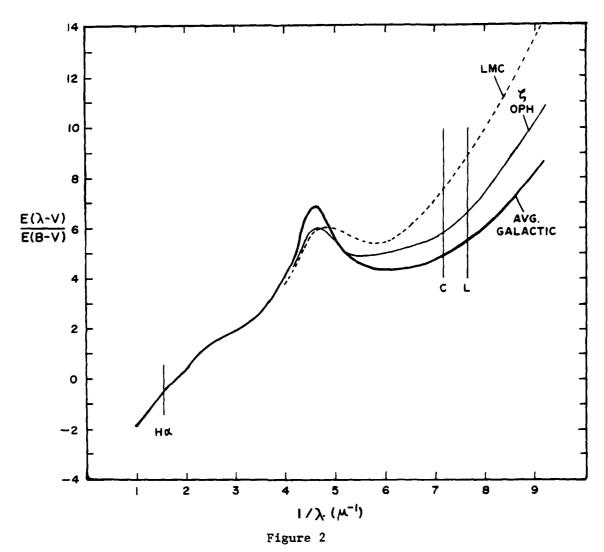
We thank Dr. Karl Henize for useful discussions. This work was supported, in part, by MASA Grant NASW-3023 to T.P.

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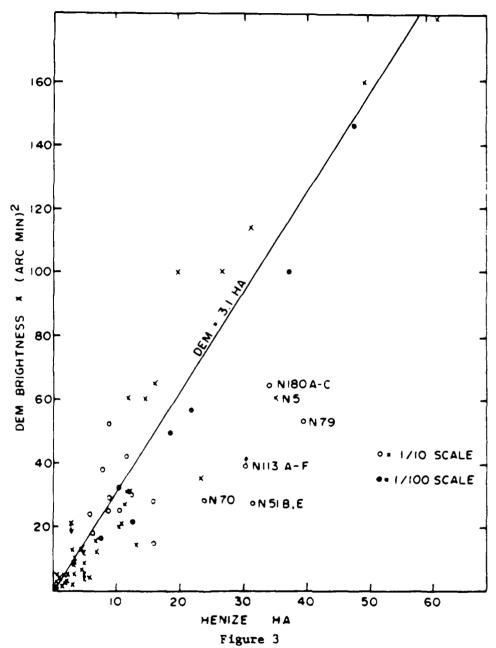
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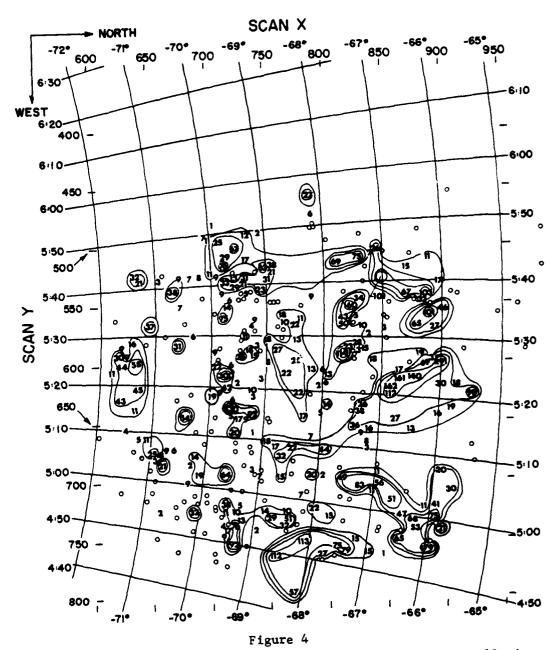
Contour plot of E(B-V) in the IMC, based on values given by Lucke (7). These were used for correcting the far-UV and H α brightnesses for interstellar extinction using the curve of Nandy et al. (3) in Fig. 2. Superimposed on the plot is an approximate RA-DEC (1950) grid, with north to the right.



Interstellar extinction curves typical of the local regions of our galaxy (8) and for the 30 Doradus region of the LMC (3). C and L indicate the effective wavelengths of the S201 imagery with CaF $_2$ corrector (1400 Å) and with LiF corrector (1300 Å).



Plot of our estimates of H α brightness x (arc min)² for emission nebulae observed by Davies et al. (1) vs. H α brightnesses of Henize et al. (11,12) for objects in common.



Contours of the Hydrogen Index (times 100) in the Large Magellanic Cloud. Contour lines are for $100 \, \text{H}$ Ind = 10, 20, 50, and 100. The vertical and horizontal axes are as for Fig. 1.

Appendix A

S201-ATLAS-LISTING TAPE AND LINEARIZED-DENSITY-MOSAIC TAPES

The listing in Appendix B of this Atlas is available on seven-track, 800-bit-per-inch, odd-parity tape. The tape was written on a Univac 1110 computer under the EXEC VIII operating system using Fortran-formatted write statements. Thus the file structure is of the Univac SDF sequential formatted record type. A more detailed description of this format can be found in the Sperry Univac 1100 Series Fortran V Library Programmer Reference (UP-7876).

There is one data file on this tape, consisting of 1796 data records of 132 field data characters each. The first data record contains a title line, "REVISED S201 FAR-UV ATLAS OF THE LARGE MAGELLANIC CLOUD." This is followed by a 132-character line of column headings, as in Appendix B (but not repeated). The meanings of the remaining 1795 data records are given in Tables Al and A2. To accommodate groups of LH objects (see text), there is a different line format for them (Table A2). Character 31 specifies the group-data-line format.

The Atlas tape file ends with a software end-of-file mark and a hardware end-of-file mark. Table A3 gives a simple Fortran program for reading the Atlas tape. The tape has been checked for errors, using this program.

Two linearized-density-mosaic tapes provide the mosaics of D_L values used in summing density volumes from frames A124 and A125 (on one tape) and A129 and A130 (on the other tape). These two tapes were written on a Univac 1108 under the EXEC II operating system, and each contains two files, one for each frame, covering the area from x=475 to x=986 and from y=381 to y=830. Each file ends with a software end-of-file mark and a hardware end-of-file mark. The simple Fortran program in Table A3 will print out the mosaics in convenient form. The mosaics are each 145 pages long; in pairs (290 pages) they require larger-than-normal storage.

Other programs can be written for listing single Lucke-Hodge objects, selecting characters 68-72 with no parentheses and no asterisk in 30, for listing Henize nebulas, selecting characters 91-98 with no parentheses, and for listing unidentified objects, selecting double minus characters 119-120 and parentheses or nothing in 68-76 and 91-97.

Table Al - Meanings of Characters in a Normal Data Line

Characters	Meaning (digits right-justified)
2-4	Frame number
6-8	x raster coordinate
10-12	y raster coordinate
10-12	•
15	Hours of right ascension
16-19	Separator (:) Minutes of right ascension rounded to tenths
21-23	5
21-23	Degrees of declination
25-26	Separator (:) Arc-minutes of declination
28-29	
1 1	x-raster interval summed (*X)
30	X (times)
31-32	y-raster interval summed (*Y)
33	An asterisk (*) indicates that area *X*Y is not rectangular
34-36	Peak density at center of image (P)
37	An asterisk (*) indicates P is not a maximum
38-40	Background density (BG)
42-46	Density volume of area summed (V)
48-49	Exposure time in minutes (E)
50	Filter type (L or C)
52-56	Density volume divided by exposure time (V/E)
57	An asterisk (*) indicates that the density exceeds 600
58-60	Reddening, RE = $E(R-V)$, in magnitudes, rounded to hundredths
61	An asterisk (*) indicates an RE value observed by Lucke
62-66	Unreddened UV flux (UF)
68-76	LH followed by one- to three-digit numbers are LH objects, parentheses mean that the Lucke-Hodge area overlaps the area summed; SAO followed by six digits means a foreground star near the area summed
76-79	North-south extent of LH object in arc-minutes rounded to tenths
81-84	East-west extent of LH objects in arc-minutes rounded to tenths
86-89	Number of blue stars in LH object (BS)
91-98	Numbers and letters of Henize nebula or nebulas (N NO.); parentheses mean that the nebula area overlaps the LH area summed; D followed by two- to three-digit numbers are DEM objects (Ref. 1)
98-103	Ha flux in units of 10 ⁻⁴ erg/s · cm ² · sterad, rounded to tenths, from Henize nebula or nebulas or DEM object (HA)
105-109	Hydrogen index, H IND. = HA/UF, rounded to hundredths
110	An asterisk (*) indicates an uncertain H IND. value; V < 10
111-120	Numbers separated by a comma or dash are NGC numbers associ- ated with the LH object; in one case the number starts with IC; in a few cases an LH number in parentheses or an
	SAO number in parentheses indicates overlaps
119-124	Six-digit number of an SAO star
125	A query (?) indicates an uncertain SAO identification; the
1	letter H (one case) indicates that the number is in the
	HD Catalog; MC followed by a two-digit number is a radio
}	source in the MC catalog (Ref. 13). If followed by SNR,
	the radio source has been identified as a supernova remnant
126-129	Visual magnitude of an SAO star, rounded to tenths (M)
131-132	Letter and digit for spectral type of an SAO star (SP)

Table A2 - Meanings of Characters in a Group Data Line

Characters	Meaning (digits right-justified)
2-4	Frame number
6-8	x raster coordinate
10-12	y raster coordinate
14	Hours or right ascension (R.A.)
15	Separator (:)
16-19	Minutes of right ascension rounded to tenths
21-23	Degrees of declination
24	Separator (:)
25-26	Arc-minutes of declination
28-30	Number of pixels summed in a group of LH objects
31	An asterisk (*) indicates a group data line
34-36	Peak density at center of group image (P)
37	An asterisk (*) indicates that P is not a maximum
38-40	Background density (BG)
42-46	Density volume in the group area summed
48-49	Exposure time in minutes
50	Filter type (L or C)
52-56	Density volume divided by exposure time (V/E)
57	An asterisk (*) indicates that the density exceeds 600
58-60	Reddening, RE = E(B-V), in magnitudes, rounded to hundredths
61	An asterisk (*) indicates an RE value observed by Lucke
62-66	Unreddened UV flux (UF)
68-77	I.H followed by one- to three-digit numbers that are separated
	by commas are LH objects in the group
79-82	Total area of the LH group in (arc-min) rounded to tenths
83	An asterisk (*) indicates a group
86-88	Total number of blue stars in the LH group (BS)
91-98	Numbers of Henize nebulas ovrelapping the LH group
111-117	NGC numbers associated with LH objects in the group

Table A3 - Simple Fortran Program for Reading the LMC Atlas and Mosaic Tapes

Line 1:		DIMENSION LINE (22)
2:		REWIND 1
3:	5	READ(1,1000,END=100) LINE
4:	1000	FORMAT (22A6)
5:		WRITE(6,1000) LINE
6:		GO TO 5
7:	100	STOP
8:		END

Appendix B

S201 ATLAS LISTING OF FAR-UV OPJECTS IN THE AREA OF THE LARGE MAGELLANIC CLOUD

The S201 Atlas listing contains 473 far-UV objects in the LMC area, each detected on one or more of the four frames: Al24 (1-min ILi exposure), Al25 (3-min ILi exposure), Al29 (10-min ICa exposure) and Al30 (30-min ICa exposure). There are 26 columns, listing data from four other catalogs, as well as S201 measurements of far-UV flux from 122 Lucke-Hodge associations [19] with associated NGC objects, from 156 Penize nebulas [17], and from 20 SAO foreground stars [24]. The column entries are defined as follows, with asterisks on column entries flagging peculiar entries as noted.

- FR. S201 Apollo frame number
- Y x coordinate in the PPS microdensitometer scan
- Y y coordinate on the PDS microdensitometer scan
- P.A. right ascension for the 1950 epoch in hours and minutes (to tenths of minutes), obtained from the LH, Henize, or SAO catalog for objects therein and from the xy coordinates for unidentified objects
- DEC. declination for the 1950 epoch in degrees and arc-minutes, obtained from the LH, Henize, or SAO catalog for objects therein and from the xy coordinates for unidentified objects
- *X number of pixels summed along the x axis, centered at X

*Y number of pixels summed along the y axis, centered at Y.

A multiplication sign between the two values *X and *Y indicates that the cataloged size of the object is matched by an area ΔxΔy in units of the area of one pixel; an asterisk indicates that the ΔxΔy area is not a rectangle but is slanted or curved. For grouped images the total number of pixels summed is listed as a single value followed by an asterisk, instead of being listed as a product of two values.

P The central (peak) density of the image, corrected for nonlinear response but not for PDS lag. An asterisk indicates that the image center (pixel at x,y) is not a density maximum.

the local background density, obtained by averaging the four density values on the centers of the four sides of the rectangle $\Delta x \Delta y$ from the mosaic of density values corrected for nonlinear response. In some images BG has a 1/2-density-unit remainder, and the listed value has been rounded upward to a whole number and is 1/2 density-unit high.

density volume = $\sum (D-BG)$ over the summed $\Delta x \Delta y$

V

V/E

RE

E, F exposure time, in minutes, and filter (L = LiF, with passband 1050 to 1600 Å; $C = CaF_2$ with passband 1250 to 1600 Å)

density volume divided by exposure, a measure of the flux reaching the S201 camera. An asterisk indicates densities > 600.

color excess in magnitudes. An asterisk indicates values measured by Lucke [20]; other values are interpolated from the contour plot, Fig. 14. Lucke's values of E(B-V) in the LMC have been increased by 0.05 magnitude for foreground reddening (Borgman et al. [10,11]).

UF density volume V/E corrected for extinction based on RE. A dash indicates a value of V/E < 0. For ILi frames UF = $(V/E)10^4 \cdot 8(RF)$. For ICa frames UF = $(V/E)10^4 \cdot 0(RE)$.

LH NO. number of association or cloud in the Lucke-Hodge catalog [14]. Numbers in parentheses are assumed to be associated with the Henize nebulas listed under N NO. or are other, overlapping LH numbers. In 23 cases, groups of two or more LH numbers are listed.

dimensions of the LH association or cloud in arc-minutes north-south (along scan x) and east-west (along scan y). The summed area \$\Delta x \Delta y\$ was generally one raster larger in each dimension to allow for the \$201 camera resolution of 3 arc-minutes. (One raster = 33 \text{ \text{µm}} on the film = 1.19 arc-minutes in the sky.) In 37 cases the area published by Lucke [7] does not agree with these dimensions, which are presumably only rough estimates. For grouped images the total area in (arc-minutes)² is listed, followed by an asterisk.

Number of blue stars (Lucke's count [20]) in the LH association or cloud

N NO. number of a nebula in the Henize catalog [17]. In many cases, the summed area ΔxΔy corresponds to several Henize nebulas; for example, 77A-E means N77A and N77B and N77C and N77D and N77E; 8, A means N8 and N8A and 26, 27 means N26 and N27. These combinations were selected after plotting the nebula positions and dimensions on a mosaic of density vs x,y. The N numbers in parentheses are near unidentified images (density maxima on two or more frames).

D MO. numbers of PEM objects, Ref. 1

HA

Fenize's $H\alpha$ intensity estimate calibrated by Dougherty, Henize, and Aller [12] in $H\alpha$ -flux units of 10^{-4} erg/s cm² sterad

- HA (cont.) summed for all nebulas listed under N NO. Their calibration was as follows: Henize "T" = 1.0 X 10^{-4} erg/s cm² sterad pixel, Henize "1" = 2.0 X 10^{-4} erg/s cm² sterad pixel, Henize "2" = 4.5 X 10^{-4} erg/s cm² sterad pixel, Henize "3" = 7.0 X 10^{-4} erg/s cm² sterad pixel, Henize "4" = 9.5 X 10^{-4} erg/s cm² sterad pixel, Henize "5" = 12.0 X 10^{-4} erg/s cm² sterad pixel. Hence, the H α intensity of N5, Henize "Int 2," dimensions 199 by 202 arc-seconds, or 2.8 X 2.8 pixels, is 4.5 X 10^{-4} (2.8 \times 2.8) = 35.2 X 10^{-4} erg/s cm² sterad. For N77A-E, the contributions of the five overlapping parts are 1.80 + 1.40 + 0.63 + 5.67 + 98.2 = 107.7, and the dimensions are 299 by 370 arc-seconds, corresponding to Δ x = 5.2 pixels and Δ y = 4.2 pixels. The summed area is 7 X 6 pixels, to allow for the S201 camera resolution.
- H IND. hydrogen index, the ratio HA/UF, or H\alpha flux per unit of unreddened far-UV flux. A dash indicates that the measured UF is zero or negative (due to measurement errors); an asterisk indicates an uncertain value because V is low.
- NGC NO. objects in Dreyer's "New General Catalogue of Nebulae and Clusters of Stars" (Mem. R.A.S. 49, Part 1, 1888) associated with LH associations or clouds. When more than two are listed by Lucke and Hodge, only the first and last are listed here.
- SAC NO. number of a star in the Smithsonian Astrophysical Observatory catalog identified with a measured image. In one case (R.A. = 5:32.2) a number from the Henry Draper Catalog is given, followed by H.
- MC NO. number of radio sources in the MC Catalog (Ref. 13). If followed by SNR, identified as supernova remnants.
- M visual magnitude from the SAO catalog
- SP spectral type from the SAO catalog

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R.A. DEC. X Y F BO Y E.F V/E RE W.55.9 GB: W.3 5X 5 75* 74 E.F V/E RE W.55.9 GB: W.3 5X 5 75* 74 E.F V/E RE W.55.9 GB: W.3 5X 5 186* B V 10C W.25 26 W.55.9 GB: W.3 5X 5 185* 142 91 30C 3 26 W.55.9 GB: W.3 5X 5 165* 17 17 3 2 26 3 3 2 26 W.55.9 GB: W.3 5X 5 16 73 31 20 16 3 3 26 16 4 26 W.56.2 70: 17 67 15 11 16 16 16 16 16 16 16 16 16 16 <th>5</th> <th>5</th> <th>354</th> <th>m ±</th> <th>32</th> <th>67</th> <th>152</th> <th>13</th> <th>0</th> <th>4</th> <th>36</th> <th>78</th> <th>- 36</th> <th>5</th> <th>m</th> <th>263</th> <th>955</th> <th>M \$</th> <th>16</th> <th></th> <th></th> <th>8400</th> <th></th> <th>959</th> <th>8 : 3</th> <th>583</th>	5	5	354	m ±	32	67	152	13	0	4	36	78	- 36	5	m	263	955	M \$	16			8400		959	8 : 3	583
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NGC NO. 5A0 NO. H					1826.35	.00 1828,35	.00 1828.35	.00 1828,35	.15 1833,37	.15 1833,37	.30 1833,37	1833.37												
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J.	108	183	9	4.7	9	58	1.7	56	197	197	86	93	6.5	3.8	12	<u>*</u>	53	Ø *	9-	9 -	11	90	57	99
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R-UV A.	5:04.6	5:04.6	5:04.6	5:04.6	9	5:04:8	Ø. ₹	5:04.8	5:04.9	6. <u>≠</u>		6. <u>₹</u>	9:04:0	6. 2	5:04:9	5:04:9	5:04:9	5:04:9	٥٠ <u>+</u>	5:04:9	5:05.1	5:05.1	1.6	5:05.1
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124 651 676 5:05.2 -70:58	58 3×	m	693	1 09	11 21	1.2	<u>*</u>	56	(1423)				191 AB	5 60	. 21				
125 652 677 5:05.2 -70:58	58 3×	m	208 2	2002	28 31	on	<u>*</u>	4	(1 H23)				191AB	7 0	. 28				
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124 785 693 5:05.2 -68:09		9×12	5	73 1004	1.	1004	. 0.5	1744						•	00.		249182	185 7	. 8 89
125 783 694 5:05.2 -68:09	* 1 X * 1 . 50		2211184	84 4012	2 31	1337	. 0.5	2323						9	00.		249185	185 7	8 89
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130 785 692 5:05.2 -68:09	-	3×15 • 667		118 22315	5 30C	7444.05	. 05	1179						0	00.		5 x 2	7 281645	7.8 89
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2 2	570	570					075	075	210	570	₹3.4	₹3.4	₹3.4	€3.4	(190)	(061)	(190)	(061)	(23)	(23)	(23)	(63)	110	7.70	080
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125 670 674 5:06.7 -70:32	10x16 324	233	1834 31	1 611	· · · · · · · · · · · · · · · · · · ·	1602	1 426	10.0 17.0	0 5	080 +	112.0		518150		
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85:89 - 6:00:5 289 68:58	68 9 x6	7.7	312	1.1	312 . 16	1828	1					•	.00 18387		
125 770 684 5:06.9 -68:28	9x 7 236	195	1208 3	31 40	403 .16	2362	:					· •	.00 18387		
X21 82:89 - 6.90 - 2.06:58 15x	12× 8 130	67 6	2152 100		215 .16	936	:). 0.	.00 18387		
-68:28	614 01X81	167 1	13800 30C		460 . 16	2007	1					0.	. 00 18387		
124 808 689 5:07.0 -67:39	ex e 11	12	1 8 1	11	18 .09	0				(21)	_	· •	.00 18467		
125 e10 689 5:07.0 -67:39	3x 2 193	181	50.3	3.	17 .09	ş.	;			(81)	<u>-</u>	•	.00 18467		
129 809 608 5:07.0 -67:39	2 X 2 57	ž.	45 10C	2	5 . 09	Ξ	,			(1.5)	=	0.	.00 18467		
130 809 687 5:07.0 -67:39	3x + 136	108	214 30C	ũ	7 . 09	91	;			1121	=	. 0	.00 18467		
124 723 676 5:07.1 -69:26	2x 2 76	7.3	- 01	11	91.01	58						•	. 00	:	
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124 551 658 5:07.2 -73:06	001 S X4	7.0	272	11.	272 .05	472							. 00	256160 6	3 A 0
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129 721 631 5:17.4 -69:38	3 × 2	081.681)01 O4	S1. * .15	<u>.</u>			1117	<u>.</u>	<u> </u>		
130 721 629 5:17,4 -69:38	e ×	630-618	195 300	21.15	27			1117	5.	.08		
124 684 627 5:17.7 -70:21	2 x 2	₩ 0B	20	11. 20 .11	67				.	00.	;	
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125 830 638 5:18.1 -67:18	5x 3	202 191	. 16	36 38 .13	- 3	1		0127	16.0	.16 1905		
129 829 639 5:18.1 -67:18	3× 4	63 50	30 1 68	£ 1 · 6 01	62	1		0127	16.0	74 1905		
130 829 635 5:18.1 -67:18	×	153 118	393 300	13 . 13	£ +	;		0127	0.91	5061905		
124 709 626 5:18.2 -69:53	X x	87 78	1 16	11. 97 .13	80+			0133	24.0	90.		
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130 778 624 5:19.3		-68:24	2 x 2	115.114	<u>.</u>	- 5	30C	-	1.0	æ				-	118	<u>.</u>	٠				
T.61:5 059 717 421		-69:38	* ** ** ** ** ** ** ** ** ** ** ** ** *		116 104	9	=	19	٠ -	622	L H+2	۶٠٥		:	(120)	0.	00	9 : 6			
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130 718 617 5:19.7		-69:38	r × r	134	734.607	1163	300	39•	394.12	111	L H4 2	5.0	6.1		(150)	0,	8161 00.	8161			
125 727 618 5	5:20.1 -6	-69:59	¥ ¥	257	232	450	31	150	<u> </u>	631	(9447)			=	(119,122)	٥.	00.	.00 1922,26	~		
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ATLAS DEC.	67:57	-67:57	-67.57	-67:57	-65:46	. 65:46	5:46	-65:46	-67:55	-67:55	-67:55	7:5	-68:41	9:4	5:22.5 -68:41	-68:41	-67:59	7:59	5:22.6 -67:59	65:29	7:52	5:22.6 -67:53	5:22.6 -67:53	-67:53	
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FAR-UR.A.	5:22.2	9.	5:22.2	5.22.2	22.	5:22.2	22.6	5:22.2	2.5	5:22.4	. se	. ea	22.5	25.	6.	20.5	5:22.6	3.6	22.E	5:22.6	22.6	25.6	22.6	22.6	22.6
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REVISED SEGI FAR-UV ATLAS FR. X Y R.A. DEC.	124 802 612	125 802	129 601 614	130 801 612	124 908 623 5:22.2	429 606 SZ1	34:59- 2:55:5 429 606 621	130 910 622	4.55:5 119 25.4	803	208 621	130 802 611 5:22.4 -67:55	124 763 608 5:22.5	765	129 763 638	130 764 606 5:22.5	798	125 800 612 5:22.6 -67:59	199	199	9 0%	808	8 0%	90	859
86 V 5	₹21	125	129	1 30	1.24	125	129	1 30	12,	125 803 613	129	1 30	12.	125 765 610 5:22.5 -68:41	129	1 30	124 798 610	125	519 661 621	130 799	124 BOY 611 5:22.6 -67:53	125 805 612	129 804 612	130 804 610 5:22.6	124 859 613 5:22.8 -66:44

REVISED S201 FAR-UV ATLAS OF FR. X Y R.A. DEC. "X 125 86C 615 5:22.8 -66:44 3X	# ¥ ₩	LARGE MAGELLANIC P BG V E.F 96*195 7 3L	CL 0UD V/E RE	jo O	И ИО.	3215	BS N NO.	7 Y	HIND. NGC NO.	SAO NO. H
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5:22.8 -66:44 3X 5:22.9 -67:12 3X	3 140•139	5 30C	0 6	n 63	;			9 9	0 1 0 0 0 .	
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601 5:23.0 -70:13 3x	TT + TT E)	0 11	0 . 15	5			130	3.6	00.	
-70:13 3x	861.961 €)	3 -11 3L	-4 . 15	5 -20			130	3.6	-, 25	
5:23.0 -70:13 3x	3 41.45	201 2-	91.0	5			130	9 · M	00.	
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-68:04 5x	(5 89 85	5 13 1L	13 . 28	8 287	, 64HJ	4.0 3.0	8 01597	S. S	.04 162128	
-68:04 5x	4228455	16 69 31	23 . 28	8 507	6 +H 7	4.0 3.0	8 01597	5.5	.02 162128	
129 797 608 5:23.1 -68:04 5X	6.5 90* BB	9 83 100	8 . 28	6 105	L H+9	4.0 3.0	16 01597	υ	.10 162128	
-68:04 5X	**5*5*5 \$. 203 30C	7 .28	8 92	6 + H 7	4.0 3.0	8 D1597	υ. υ	.11 102128	
-66:25 3x	63 79* 79	11 0 8	0 . 10	0 0			9	5.0	00.	HC 36
-66:25 3x	(3 202 201	16 31	2 .10	9 0			9	. D	1.05	HC 36
-66:25 3X	(3 62• 60	9 6 100	1 13	9			3	0.0	3.15	MC 36
-66:25 3x	3 167•161	1 24 30C	0:.1	0 2			9	5.0	3.15	MC 36
-71:38 8x	17 82 TT	ו 19 ור	91. 61		111 (SA0256180?	1 2 0	198	58.3	.76	
-71:38 8x	x 7 212 190	3 284 3L	95 . 16		556 (SA0256180?)	0 7 1	961	58.3	51.	

REVISED SZOT FAR-UV ATLAS FR. X Y R.A. DEC.	0F THE	LAHGE 1	LANGE MAGELLANIC P BG V E.F	NIC CLOUD E.F V/E	UD E RE	UF	1215 . ON HJ	.0N N 28	4	HIND. NGC	0	SA0 NO. M 5
129 624 601 5:23.2 -71:38	8x 7	31. 44	¥15	100	31 . 15	6	(SA0256180?)	96-	58.3	£6.		
130 626 599 5:23.2 -71:38	8x 7 1	196 104	160	30C	25 . 16	601	109 (540256180?)	961	58.3	. 11		
124 625 600 5:23,2 -71:37	5x 3	92 74	99	=	68 . 05				0.	00	•	25618077.4 A3
125 627 603 5:23.2 -71:37	5x 4 2	212 187	2 36	31	79 . 05	137			9.	00.	•	256180?7.4 A3
129 626 601 5:23.2 -71:37	5 X S	35 36	348	100	35 . 05	55			0.	00.	-	25618077.4 A3
130 626 599 5:23.2 -71:37	9 x 6	196 85	0681	300	63 . 05	Ø.			ø.	00.	•	25618077.4 A3
124 707 601 5:23.3 -69:54	2× 3	82. 82	7	=	714	32		131	2.0	60		
125 707 602 5:23.3 -69.54	2 x 3 2	212.211	<u>*</u>	31		6.3		131	٥٠٤	. 12		
129 707 602 5:23.3 -69:54	ex 3	63. 67	0	100	* -	m		131	8.0	. 92		
130 708 600 5:23,3 -69:54	2x 3 1	176•186	on \$	300	¥ .	7		13:	0 · d	65 .		
124 788 602 5:23.5 -68:13	2 X 2	84 78	61	11	19 . 20	173		0110	10.0	60.		;
125 788 607 5:23,5 -68:13	2 x 2 2	205 194	36	31	12 .20	601		0110	10.0	.15		;
129 788 607 5:23.5 -68:13	2 x 2	9, 19	- 69	100	6 . 20	3.7		0110	0 . 0 1	£ +		;
130 789 604 5:23.5 -68:13	5x 4	162 111	477	300	16 .20	100		0110	10.0	9 .		1
124 862 610 5:23 6 -66:41	3 × £	77 08	28	1:	28 . 09	67		0182	13.3	\$.		
125 863 611 5:23.6 -66:41	5 X 5	210 197	÷ 5	31	48 . 08	116		0182	133	*		
129 863 612 5:23.6 -66:41	e x e	71 50	2 4 6	100	25 . 08	58		0182	13.3	.31		
130 864 611 5:23.6 -66:41	1 7 X 7	187 121	*	300	47 . 08	96		D182	13.3	9 1		
124 637 600 5:24.0 -71:23	13x 8	80 - 78	59	11	59 . 17	386	(1,450)	199.200	0 121 0	9		
125 639 600 5:24.0 -71:23	13x 8 2	212 192	270	31	11. 06	583	11 H501	199.200	0 151 0	3.0		
129 637 600 5:24.0 -71:23	13× 8	75 49	691	100	17 . 17	6	(LH50)	199.200	0 121.0	2.21		
130 638 595 5:24.0 -71:23	1 3x B 1	151112	1175	300	39 .17	186	11.450)	199.200	0 121 0	96		
04-318-601-5:24.1 -69:40	5x 7	93 - 88	,	=	¥1 13	172		138A-J	*	60		MC 395NR
125 713 601 5:24.1 -69:40	5x 7 2	248.234	177	31	59 .13	2 + 8		132A-J	* = = = = = = = = = = = = = = = = = = =	90		MC 395NR
129 720 600 5:24.1 -69:40	5x 7 139	39 95	356) O C	36 . 13	1.9		132A-J	* · · · · · · · · · · · · · · · · · · ·	£ -		HC 395NR

. ₹	A. DEC.	•	1			,		RE		LH NO.	3215	į.	8	N NO.	H AH	HIND. NGC NO.	SAO NO.	I
130 721 598 5:24.1 -69:40	1.1 -69:40	- 5X	¢ 65	567	1595 3	300	E G	<u>m</u>	175					136A-J	*	60.	MC 395NR	
124 752 598 5:24.4	. 4 - 68:58	2 x 3	a	69	,	=	•	01.	5 1					137AB	ē.	.05		
125 753 600 5:24.4	.4.4 -68:58	2 x 3	210	1.207	ø	¥	٠. م	01.	9					137AB	ń	. 10		
129 752 599 5:24.4	95:89- 4.4	2× 3	59.	98	-	301	0	01.	0					137AB	ĸ,	.00		
130 753 597 5:24.4	1.4 -68:58	e x		**!*6*!	E E	30C	•	0.1	9					137AB	ĸ,	00.		
124 637 598 5:24.5	.5 -71:23	6 × 8	893	77	3.7	=	37	11.	245	, 450	7.0	0.8	*	(500)	a.	0.0		
125 639 597 5:24.5	1.5 -71:23	6 × 8		061.961	4.7	¥	91	. 17	70 -	. 05H7	7.0	8.0	<u>*</u>	(500)	0.	. 30		
129 038 599 5:24.5	1.5 -71:23	6 × 8	16	20	147	1 0 0		. 17	1.	, usu	7.0	8.0	<u>*</u>	(200)	0.	00.		
130 438 597 5:24.5	1.5 -71:23	6 × 8	212	- 63	706 3	30c	 **	.17	·	LH50	7.0	8.0	<u>*</u>	(500)	a .	. 00		
124 773 598 5:24.8	.8 -68:33	7 x 8	6	3	69	7	69	£ 1.	290					138.A-D 6	÷. +9	.30		
125 777 595 5:24.8	6 - 68:33	7 x 8		112.822	302	31	101	. E.	424					138.A-D 6	÷ .	.20		
129 774 598 5:24.8	8 -68:33	7 × 8	106	99	530	1 0 0		E .	175					138.A-D 6	÷ .	. 50		
130 775 596 5:24.8	.8 -68:33	7 x B	320	167	1653 3	30C	55	E -	182					138.A-0	# · # 9	6 0		
125 711 599 5:24.9	. 9 - 69:53	3x	556	215	102	31	 *E	- - - -	159					(131)	0.	. 00		
129 709 598 5:24.9	9 -69:53	r × s	68	69	358	100	36	<u>*</u>	1 30					(131)	ø.	. 00		
130 710 594 5:24.9	.9 -69:53	6 X 9	248	001	1831 3	30C	- - -	*	125					(131)	0	.00		
124 879 604 5:25.4	.4 -66:23	15×12	90	60 19	329	=	3. 658	. 07	713	(1, 453)				48.A-E 27	270.3	ž.		
125 881 604 5:25.4	.4 - 56:23	15x12		205.204	609	×	P 03	. 07	0 5 5	(1,453)				48.A-E 27	270.3	.72		
129 882 604 5:25.4	.4 -66:23	15x12	19•	9	101	100	0.	. 07	192	(1,453)				48.A-E 27	270.3	1.65		
130 883 602 5:25.4	.4 -66:23	15x12	236	236.159	3354 30	300	112 .0	. 07	213	(1,453)				48.A-E 27	270.3	64.1		
124 822 595 5:25.5	.5 -67:30	ν ×	9	901	101	=	. 37 . 11		360 1	1 121	٠.	3.0	ស	(55)	G.	00		
125 823 597 5:25.5	.5 -67:30	t.	333	333*30%	÷	31	161 .11		543 (1 1511	٠. 5	3.0	5	(52)	0.	.00		
129 824 596 5:25.5	.5 -67:30	ς.	4 2 2	6 * 9	94-1	100	. 69	=	1 061	1 1511	5. 5	3.0	r.	(55)	o	00.		
130 824 594 5:25.5	.5 -67:30	¥ £	956	667	1354 30	30c	451.1.		1 521	1 1511	5.1	3.0	r	(55)	o .	00.		
124 630 591 5:25.7	.7 -71:32	3× 3	9 2	7.5	'n	=	 m	.17	32				•	201,202	19	. 0 3		

REVISED 5201 FAR-UV ATLAS OF THE LANGE MAGELLANIC CLOUD

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SAO NO. H				MC+0.415NR	MC40.415NR	MC40.415NR	MC+0.415NR					1	;	1	:	;									
HIND. NGC NO.	57.	00.	. 00	8461 85	.75 1948	8461 56	8461 61 1	8461 00.	8+61 00	8+61 00	8461 00	<u>*</u>	=	1.15	. 74	00.	. 16	.17	16.	.65	= :	. 10	. 34	* 5.	.00 1955
										•	•					0									9
Ĭ				20.3	20.3	20.3	20.3	•	•	•	•	7.0	7.0	7.0	7.0	•	¥ . 0 .	* 0	* O -	10.4	1.17.7	4.7	r. 7. 1	1 47.7	
02 2	201,202	201.202	201.202	4 BA - C	48A-C	484-C	484-C	(84)	(84)	60 + -	6843	4610	₩610	*610	+610		145	5 + -	2 7 -	1+5	140.143	140.143	140.143	140.143	(55)
98								1.5	1.5	15	S ~														1.7
3715								4.5 4.5	F. 5 F. 5	4.5 4.5	# . P.														392.20
LH NO.				(1, 452)	(1,452)	(1,452)	(1, 452)	+ 25H1	LH52	LH52	t 452 +						(1 446)	(P + P)	(1,446)	(11.1146)	(1,457)	(1,457)	(1 H57)	(LH57)	725 LH51.54
, J	ø	٥	0	-	32	4	20	951	169	55	55	7.2	86	ō	<u>:</u>	4 68	98	19	5	2.	539	593	100	256	725 L
RE	. 17	. 17	.11	. 07	. 07	. 07	. 07	. 07	. 07	. 07	. 07	113	.17	. 17	. 17	51.	. 1.2	51.	. 12	٠.	=	Ξ	=	=	=
3//E	-	0	٥	5	5	1.3	Ξ	72	7.8	69	69	Ξ	1.5	~ i	₩	6	23	15	KD.	7	11.5 211	176	99	8	215
	31	100	300	=	ĸ	1 0 C	30¢	=	31	100	300	1	31	100	30C	31	=	31	J 0 C	30C	1	3	100	30C	1
MAGELLANIC 5 V E.F	*	*	'n	6 -	ş Q	125	3 5 M	12	233	286	957	=	9	\$ 2	75	278	63	63	50	202	160	527	657	2600	215
OF THE LARGE MA	3x 3 193+193	3× 3 +1 + +0	3x 3 94 93	6x 3 90 86	6x 3 228 217	6x 3 1u1 • 86	6x 3 310 242	6x 6 90 83	6x 6 224+211	6x 6 104 79	6x 6 300+214	PK TT S XS	38 1 96 1 8 XE	2x 2 43 37	2x 2 112 89	5x 5 238 219	5x + 88 86	5x + 23++233	5x 4 103 92	5x + 290 2+0	6x 6 91 87	8x 8 243+232	9x 8 100 • 67	8x 8 308+263	37. 114.103
ATLAS DEC.	11:32	-71:32	-71:32	-66:19	-66:19	-66:19	-66:19	-66:17	-66:17	-66:17	-66:17	-71:40	-71:40	04:16-	-71:40	-69:59	-69:28	-69:28	-69:59	-69:58	-69:15	-69:15	-69:15	-69:15	-67:30
REVISED S201 FAR-UV ATLAS FR. X Y R.A. DEC.	125 631 592 5:25.7 -71:32			5.7 -(5.7 -1																
SI FAR-	5. 5.	591 5:25.7	5:5	5: 2	5:25.7	5:5	5.		603 5:25.7	603 5:25.7	5:5		591 5:25.8	593 5:25.8	591 5:25.8	5:5	2:5	5:25.9		591 5:25.9	5:2	5:25.9	5:25.9	5:2:	5:8
52 0 Y	1 596		1 586	9 601	9 9	1 603	5 601	109 9	\$ 603	5 60	6 601	3 587		593		9 578	9 592	9 592	9 593		165 8	9 592	592	9 589	2 593
REVISES FR. X	25.63	129 631	130 631 588 5:25.7	124 885 601 5:25.7	125 886 603	129 884 603 5:25.7	130 885 601 5:25.7	124 886 601 5:25.7	125 085	129 885	130 886 601 5:25.7	124 623 587 5:25.8	125 621	129 621	130 625	125 729 578 5:25.9	124 729 592 5:25.9	125 729 592	129 729 593 5:25.9	130 730	124 738 591 5:25.9	125 739 592	129 738 592	130 739 589 5:25.9	124 822 593 5:25.9
æĮ	2	2	-	2	=	Ξ	-		2	7	-	=		_	-	,		,	Ž	-	,,,	7	-	-	9

R P F P	5£ D 520	REVISED SZOJ FAR-UV ATLAS FR X Y R.A. DEC.	V ATLAS DEC.	0F THE	P BG V E.F	AGELLA V		3 / A 0 0 0 D	æ	Ų	LH NO.	3215		X 58	0	¥	ND.	HIND. NGC NO. SAO NO	SAO NO.	Ε
125 82	965 828	6.5:25.9	67:30	-62	3+8+302	470	ž	157	Ξ	529	529 LH51.54	392.2		17 ((52)	•	00	1955		
129 824	954 286	6.55.25.9	-67:30	37.	422 187	2018	100	202	=	556	556 LH51,54	392.24		17 ((55)	•	00.	1955		
1 30	824 592	9:55.9	-67:30	37.	790.562	3989	30C	1334.11	Ξ	366	366 LH51,54	388.8		17 ((52)	•	00.	1955		
154	888 602	8 5:25.9	-66:14	16x 7	86 . 82	34	1	3	. 07•	73	73 LH52.53	0.61	6.0	34 0191	_	5	0.0	8461 80		
125 666	209 888	2 5:25.9	-66:14	1 8x 7	2151205	450	3	150	. 07•	325	325 LH52.53	0.61	6.03	34 0191	_	Ø.	. 02	9 4 6 1		
129 887	887 602	8.55.5	-66:14	18X 7	99 • 06	721) O C	7.2	.07•	137	137 1452.53	19.0	0.0	34 0191	_	ۍ	5 0.	9461 40.		
130 867	009 490	0 5:25.9	-66:14	18x 7	253.164	3048	30c	102	.07•	194	194 LH52.53 19.0		6.0	34 0191	_	5.0	.03	8461 20		
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5201 FAR-UV Y R.A.	5:32.0	5:32.0	32.0	5:32.0	5:32.0	32.0	32.0	32.0	5:32.0	5:32.1	1 . 2	5:32.1	52.1	567 5:32.1	5:32.1	5:32.1	5:32.1	5:32.1	5:32.1	1 . 5	5:32.1	5:32.2	5:32.2	5:32.2	5.2
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REVISED FR. X	173	130 774	853	129 823	823	628	831	830		764	125 765	766	130 765	124 815	ž	129 815	918	124 877	878	129 877	130 678	709	710	709	709
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N V	9	8,1	9 * 1 >	0,	55.A	55.A	55.4	55.A	0232	0232	0232	0232	9		9,1,	9	57.A-E	57.A-E	57.A-E	57,A-E	1231	(57)	(57)	(57)	(28)
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REVISED SZOI FAR-UV ATLAS FR. X Y R.A. DEC.	5:32.2	5:32.2	5:32.2	5:32.2	5:32.3	5:32.3	5:32.3	5:32.3	5:32.4	5:32.4	5:32.4	5:32.4	5:32.4	5:32.4	5:32.4	5:32.4	5:32.5	5:32.5	5:32:5	5:32.5	5:32.6	5:32.6 -	5:32.6 -	5:32.6 -	5: 32.6 -67: 32
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SAO NO												HC 56	HC 56	MC 56	MC 56										
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NGC NO	(LH78.79	ILH78.79	11 H78.791					2021	2021	1202	1202					.00 2002-34	2002	.00 2002-34	.00 2002-34	2025	2025	2025	. 60 2025		
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₹ ⁻									9							9- (9-							- 6
SZOI FAR-UV ATLAS Y R.A. DEC.	5:32.6	5:32.6	5:32.6	5:32.7	5:32.7	5:32.7	5:32.7	5:32.8	5:32.8 -67:31	5:32.8	5:32.8	5:33.0	5:33.0	5:33.0	5:33.0	5:33.0 -66:56	5:33.0	5:33.0 -66:56	5:33.0	5:33.1	5:33.1	5:33.1	5:33.1	5:33.3	125 713 561 5:33.3 -69:48
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SE 0 ×				285		585	285		552	325	952	971	671	611	671	158		351	352	320	320		322	113	113
REVISED FR. X	125 824	129 824	130 824	124 682	125 682	129 682 564	130 682	124 854	125 625 562	129 825	130 825	124 778 561	195 779 561	129 779 562	130 779 560	124 851 562	125 851	129 821	130 852	295 029 421	125 620 562	229 621	130 622	E11 42.	125 7

HA HIND. NGC NO. SAO NO. M	1.9 .00	00 . 9.1	.00 .00 .00	.0 .00 2021	.0 .00 2021	.0 .00 2021	.0 .00 2002-34	.0 .00 2002-34	.0 .00 2002-34	.0 .00 2002 - 34	.0 .00 (LH75,79)	.0 .00 (LH75,79)	.0 .00 (LH75,79)	.0 .00 (LH75,79)	.0 .00 2021	.0 .00 2021	. 0 . 00 2021	. 0 . 00 2021	2.7 .05	P.7 .20	2.7 .36	۶.٦ .00	.00 -00 -00 -0 -0 -0 -0 -0 -0 -0 -0 -0 -0	.0 .00 2028	.0 .00 2028
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FAR-UV R.A.	3.3	5:33.3	5:33.3	5:33.3	5:33.3	5:33.3	5:33.3	5:33.3	5:33.3	5:33.3	5:33.6	5:33.6	5:33.6	5:33.6	5:33.7	13.7	5:33.7	3.7	5:34.0	5:34.0	5:34.0	5:34.0	5:34.1	5:34.1	-
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130 780	80 548	9 5:35.4		-68:58	5 x 6	156	113	650	30C	22	. 25	220	(1,485)					Ö		.00 20427		
124 7	124 717 550	0 5:35.5		-69:44	14X17 112	118	3	1378	=	1378	. 30	37953	(LH81.87			-	154.AB 12	1288.4	. 07	7 2033.48	HC61.65.67.7	7.
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1 30 7	130 718 549	9 5:35.5		1 44:69-	14×17	128	179	30467	30C	10164.30 16102	. 30	5019	(LM81.87	2		-	154.AB 1288.4	5 8 8 · t	91.	5 2033.48	HC61.65.67.71	7
154 851	155 126	1 5:35.6		-67:35	0 1 x 8	9	69	152	=	152	Ξ	512	(LH82,88	â		ñ	56.59A-C	543.2	1.37	04-6202 /	HC 64	
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124 751 400 6:09.1 -68:50 12x11 399 86 4	4918 1L	4918 .05	9246					•	00.		248461 5	68 2
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124 520 418 6:15.9 -73:36 6x 6 153 77	11 506	50. 506 1	1572					0.	00.		256286 6	68 89
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129 522 418 6:15.9 -73:36 11X12 324 27 7	7033 10C	20. 201 05	*					•	00		256286 6	68 89
130 522 416 6:15.9 -73:36 14X14 671 67 33	33900 300	20.0511	1790					0	00.		256286 6	68 83
124 589 390 6:19.1 -72:07 3X 4 88 82	15 54	50.54 1	72					0.	00.		256290 8	0 A 0 .
125 590 391 6:19.1 -72:07 4x 5 233 210	227 31	50.92	132					0.	00.		256290 8	0 A 0 .
129 590 391 6:19.1 -72:07 5x 5 74 28	451 10C	50.64	1.1					•	00		256290 8	8.0 A0

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E UF LH NO.	5 8 5		1 - 2 1	0 01.0	0 01.0
CLOUB V/E RE	34 . 05	01.	7 .10		0 . 10
AGELLANIC V E.F	1610 30C	<u>-</u>	22 31	0 10C	-6 30C
REVISED SZOI FAR-UV ATLAS OF THE LARGE MAGELLANIC CLOUD FR. X Y R.A. DEC. «X «Y P BG V E.F V/E RE UF LHNO.	1 -72:07 6x 6 20v 73 1610 30C 54 .05 85	5 -71:35 2x 2 86 85 4 1L 4 10 12	5 -71:35 EX 2 224+232 22 31 7 .10 21	5 -71:35 2x 2 30+ 30	5 -71:35 2x 2 75 76 -6 30C
REVISED SZOL FAR-UV A FR. X Y R.A.	130 591 388 6:19.1 -72:07	124 612 383 6:19.5 -71:35	125 612 385 6:19.5 -71:35	129 613 387 6:19.5 -71:35	130 613 385 6:19.5 -71:35

BBRKPT PRINTS

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SUPPLEMENTARY

INFORMATION

ERRATA - NRL MEMORANDUM REPORT 4660

REVISED LISTING - S201 FAR-ULTRAVIOLET ATLAS OF THE

LARGE MAGELLANIC CLOUD

Please replace pages 5 through 8 in your copy of this document with the attached corrected versions (equations 2 through 6 of the original are incorrect). Also, in the paper "Distributions of Hot Stars and Hydrogen in the Large Magellanic Cloud" by T. Page and G. R. Carruthers (Astrophysical Journal, 248, 906, 1981) equations (2) through (6) should be corrected as indicated herein. The next to last paragraph on page 7, as corrected in the attached, also applies to the discussions of UF and H Ind in that Reference and its contained Table 1.

at over 100 places in the LMC. This index was first presented as a rough measure of the hydrogen near hot stars or star groups detected on our far-UV images. That is, if the ionizing extreme-UV (λ < 912 Å) flux is assumed roughly proportional to the far-UV flux, then the intensity of H $_{\alpha}$ emission is related to the local hydrogen density. Here, we present a revised determination of H Ind and its variation over the LMC, using a more recent determination of the IMC extinction law, allowing for extinction at H $_{\alpha}$ as well as in the UV, and utilizing additional data on the H $_{\alpha}$ brightness distribution in the IMC.

The far-UV flux values are proportional to the measured density volume, V (corrected for nonlinearities and background) divided by the exposure time, E, in minutes. As shown in the <u>Revised S201 Catalog of Far-UV Objects (NRL Report 8487)</u>, a density-volume

$$V = 0.037 \text{ n}$$
 (1)

where n is the number of photoelectrons forming the far-UV image. Thus,

$$V/E = 2.22 \text{ n per sec}$$
 (2)

where E is the exposure time in min, and n/sec is related to the photons arriving each sec from the object. The detection efficiency (photoelectrons per photon, based on preflight calibrations) of the S201 Camera in the imaging mode averages 0.05 over the range 1050-1600~Å with the LiF corrector, and 0.04 over the range 1250-1600~Å with the CaF₂ corrector. Hence, the photon flux in these wavelengths is

$$N_L = n_L/0.05(30.0) = 0.300 (V_L/E) \text{ photons/sec cm}^2 \text{ for } 1300 \text{ } \frac{R}{+} 250 \text{ } \frac{R}{+},$$
 (3)

 $N_C = n_C/0.04(30.0) = 0.375 \ (V_C/E) \ photons/sec \ cm^2 \ for 1400 \ Å <math>\pm 150 \ Å$, where 30.0 cm² is the aperture area of the S201 camera. Since these photons each carry 1.52 x 10^{-11} erg and 1.42 x 10^{-11} erg respectively, the far-UV flux is

$$F_L = 4.92 \times 10^{-10} \ (V_L/E) \ erg \ sec^{-1} \ cm^{-2}$$
 (5)

$$F_C = 5.33 \times 10^{-10} (V_C/E) \text{ erg sec}^{-1} \text{ cm}^{-2}$$
 (6)

These were corrected for interstellar extinction, based on previous estimates (7) of the visual reddening (RE = E(B-V)). In order to estimate reddening for all our measurements of V/E, for which specific values of RE were not available, we plotted Lucke's (7) RE values and sketched in

contour lines (see Fig. 1). Although Lucke's 81 measured values are good to \pm 0.05, corresponding to \pm 16 to \pm 17% in corrected ultraviolet flux, UF, there is inevitably some uncertainty in the interpolated values of RE, due to small scale variations in the extinction at a given distance, and the uncertainty in distance to an object along the line of sight. The stellar associations for which Lucke determined RE may lie in front of or behind far-UV objects with nearly the same celestial coordinates. However, it is highly likely that an LH cluster and an associated Henize nebula are in close 3-dimensional proximity.

In the Atlas, we used the "average" galactic interstellar extinction curve of Bless and Savage (8). However, measurements with the ANS satellite (9,10) in the 30 Doradus region, and with IUE (3) there and elsewhere in the LMC indicate a higher ratio of far-UV extinction to E(B-V) in the LMC than is typical in the local region of our galaxy (see Figure 2). Using the extinction curve of Ref. (3) with $A_{\lambda} = 3$ E(B-V) + E(λ -V), we have, for effective wavelengths of 1300 Å (LiF corrector) and 1400 Å (CaF₂ corrector), E(1300-V)/E(B-V) = 8.97 and E(1400-V)/E(B-V) = 7.09. Therefore, the ultraviolet fluxes corrected for reddening are

$$UF_{L} = F_{L} 10^{4.8} RE$$
 (7)
 $UF_{C} = F_{C} 10^{4.0} RE$ (8)

As expected, UF_L values for an object are generally larger than the UF_C values because of the wider bandpass and larger extinction correction at the effective wavelength of 1300 Å. The scatter in the LMC extinction curve of Nandy et al. (3) is about 0.2 mag. The extinction correction at H α is assumed to be $A_{6563} = 2.5$ RE; hence the corrected H α flux is UHA = HA·10^{RE}, approximately, where HA is the H α flux as measured by Henize et al. (11,12) in units of 10^{-4} erg/cm² sec sterad. The HA values given here are often summed for several close H II regions that could not be separately resolved on our S201 photos. For instance, N180A-C means the summed flux from N180A, N180B, and N180C. In order to get a single hydrogen index representing all measurements of a given object, we averaged the values for two ILi frames with 1/2 the values for two ICa frames:

$$H Ind_{L} = UHA/UF_{L}$$
 (9)

$$H \ Ind_{C} = UHA/UF_{C} \tag{10}$$

H Ind = (H Ind_{L1} + H Ind_{L2} +
$$1/2$$
 H Ind_{C1} + $1/2$ H Ind_{C2})/4 (11)

The major errors in V/E, UF, and H Ind are due to uncertainty in background, b. As can be seen from the isodensity contour plots in the Atlas, many of the objects measured are in regions where the background density is changing. The local background was estimated on mosaics of d, taking the first minimum in d in each of four directions from the peak density, along +x, +y, -x, and -y, and averaging these to get b. The background is high and posed the most difficulties on the 3-min ILi exposure, frame A125.

The HA values are probably good to \pm 10%, although values near zero are subject to larger percentage errors. In fact, DFM, in a careful survey of a 5-hour exposure with the SRC 48-inch Schmidt camera using an interference filter with 100 Å bandpass centered on E_{α} and [NII], found the faint Henize H II regions much larger, and detected 100 more, most of them fainter than Henize's limit. They give no quantitative measurements of brightness, but use the steps vf (very faint), f (faint), fb (fairly bright), b (bright), and vb (very bright). We calibrated this scale against HA by assigning the numbers vf = 1, f = 2, fb = 3, b = 5, vb = 10, and multiplying by the dimensions given in arc-min. For instance, a faint (f) nebula of size 3.5' x 2' has a brightness (arc-min)² of 2 x 3.5 x 2 = 14. Fig. 3 is a plot of these values against HA for 64 cases where the DEM dimensions are roughly the same as Henize's. To a fairly good approximation,

DEM brightness $(arc-min)^2 = 3 \text{ HA}.$

(12)

Using this calibration, we could fill in 227 H II regions at positions in the LMC where we had measured far-UV flux, leaving out only 19 DEM objects of the total of 356. (These positions were all searched on our mosaics.)

In the Revised Listing (Appendix F), we list density volumes for 473 objects or regions in the IMC. We also list values of UF, defined here as simply the density volumes corrected for extinction as per Equations 7 and 8. (True UF values, in ergs/sec cm², can be obtained by multiplying by the factors 4.92 x 10^{-10} for F_L and 5.33 x 10^{-10} for F_C , respectively.) Likewise, the H Ind values are the corrected density volumes divided by UHA. Figure 4 is a contour plot of H Ind (times 100), individual values of which are given in the Listing.

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